

Overview continued

How atomic/molecular structure \rightsquigarrow
what you see in e.m. spectrum.

1st approximation to atomic structure
= H-like ions, non-relativistic, no spin
means?

H itself is of special importance because

- H is most abundant element in Universe, 75% by mass, 90% by number of nuclei (But atomic matter makes up only 4% of total mass-energy of Universe. How do we know?)
- H lines prominent in astron spectra.
- H plays major role in ionization/thermal equilibrium
 - Sun & stars
 - Gas planets
 - Interplanetary / interstellar / intergalactic medium
 - Star formation
 - Cosmic Microwave Background

Atomic physics of H-like ions is well understood theoretically, as is He-like ions.

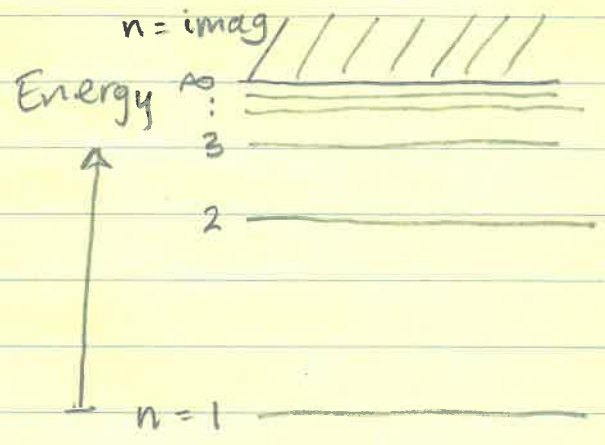
Atom physics of multi-electron ions is in progressively worse theoretical shape

- rely more on experimental data, especially for neutrals, which are easiest to prepare in lab.

Energy levels of H

$$E_n = -\frac{1}{2n^2} \text{ a.u.}$$

$$= -\frac{1}{n^2} \text{ Rydberg}$$



Each energy level is $2n^2$ -fold degenerate

↑ 2 electron spin states ↗ various angular momentum states

This ignores nuclear spin (depends on nucleus).

Each individual state is an eigenstate of mutually commuting operators:

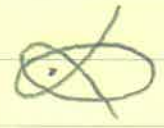
Quantity	Operator	Quantum number	Values
Energy	Hamiltonian H	n	1, 2, ..., infinity
$ Ang\ mom ^2$	L^2	l	0, 1, ..., n-1
Ang. mom z-cpt	L_z	m	-l, ..., l

m (magnetic) states degenerate because of isotropy. ^{what?}
why?

l states degenerate because of accidental symmetry in 2-body Coulomb problem that causes classical orbit to close



closed orbit



general non-closed orbit

Relativity splits l-state degeneracy.

Selection rules - Allowed transitions

Not all radiative transitions between energy levels are equally fast.

Fastest are those involving emission/absorption of photon with

(i) unit ang. mom \hbar = "dipole" photon

(ii) odd parity
= "allowed" or "permitted" transition.

Parity = behavior of wavefunction
under coordinate inversion

$$P\psi(\vec{x}) = \psi(-\vec{x}) = \pm \psi(\vec{x})$$

\uparrow
+ = even parity
- = odd parity.

Total parity is rigorously conserved in any reaction, eg



Almost absorption lines in stellar spectra are from allowed transitions.

BUT in astronomy. Forbidden lines are important in low density environments means?

such as interstellar / interplanetary space.

Ex 1 / $[O III]$ 5007 Å "nebulium" green line
 ↑ ↑
 brackets denotes "forbidden"

$O III$ or O^{++}

$Z = 8$, 6 electrons

has 2 valence electrons with $l = 1$.

Ground electronic configuration $/// // /$

split into 3 ($L = 0, 1, 2$)

by electrostatic interaction



$$\text{Parity} = (-)^{\sum_{\text{elec}} l}$$

= same for all levels of ground elec config

⇒ transitions between levels of ground elec config are all parity forbidden.

Such transitions don't happen in lab, why not?

Ex 2 / [OI] 6300, 5577 auroral lines
 pink green
 1000s k
 decay decay

Prob Set.

Molecular structure overview

3 ~ independent modes.

(a) electronic ~ like atoms

$$\Delta E_{elec} \sim \text{a.u.}$$

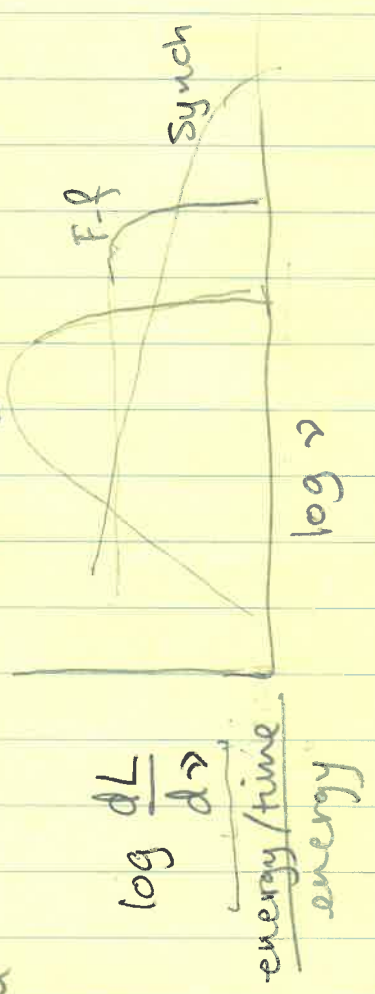
(b) vibrational ~ simple harmonic oscillator

$$\Delta E_{vib} \sim \left(\frac{m_e}{m_x} \right)^{1/2} \text{ a.u.}$$

(c) rotational ~ solution of rotator

$$\Delta E_{rot} \sim \left(\frac{m_e}{m_x} \right) \text{ a.u.}$$

	Radio	Microwave Infrared	Visible	Ultraviolet	X-ray	γ-ray
Atomic	Hyperfine Radio rec. ... recomb ...	Fine structure	Elec split Forbidden Balmer H Lyman H	Allowed	High-Z allowed High-Z H-like	(nuclear)
Molecular	Rotational	Vibrational	Electronic	(molecules dead)		
Continuum						
Blackbody = Planck	CMB 2.7K	Dust 20-40K	Planets Stars 300K -10,000K	Hot WDs 10 ⁵ K	Neutron stars 10 ⁶ K	
Free-free = Bremsstrahlung = "Thermal"		→	extending up to	$h\nu \approx kT$		
Synchrotron = "Non-thermal"		→	extending up to	$h\nu \approx h\nu_{\text{gyro}} \gamma^3$		$\gamma = \text{Lorentz factor}$



why?